

Background and imaging simulations for the hard X-ray camera of the *MIRAX* mission

M. Castro,¹★ J. Braga,¹★ A. Penacchioni,¹ F. D’Amico¹ and R. Sacahui^{1,2}

¹Instituto Nacional de Pesquisas Espaciais, Av. dos Astronautas 1758, CEP: 12227-010, São José dos Campos, SP, Brasil

²Escuela de Ciencias Físicas y Matemáticas, Universidad de San Carlos de Guatemala, Ciudad Universitaria, zona 12, Guatemala

Accepted 2016 March 25. Received 2016 March 2; in original form 2015 December 8

ABSTRACT

We report the results of detailed Monte Carlo simulations of the performance expected both at balloon altitudes and at the probable satellite orbit of a hard X-ray coded-aperture camera being developed for the *Monitor e Imageador de RAios X (MIRAX)* mission. Based on a thorough mass model of the instrument and detailed specifications of the spectra and angular dependence of the various relevant radiation fields at both the stratospheric and orbital environments, we have used the well-known package *GEANT4* to simulate the instrumental background of the camera. We also show simulated images of source fields to be observed and calculated the detailed sensitivity of the instrument in both situations. The results reported here are especially important to researchers in this field considering that we provide important information, not easily found in the literature, on how to prepare input files and calculate crucial instrumental parameters to perform *GEANT4* simulations for high-energy astrophysics space experiments.

Key words: atmospheric effects – balloons – instrumentation: detectors – methods: numerical – space vehicles: instruments – techniques: image processing.

1 INTRODUCTION

Estimation of the energy spectrum of the background signal and its spatial distribution over the detector plane is crucial for the design of hard X-ray and γ -ray astronomy telescopes. In the case of an observation of a point source from an orbital space platform (i.e. a satellite), the background consists in the diffuse electromagnetic radiation coming through the telescope aperture, emission from other sources in the field of view (FoV) and the instrumental background, which arises from interaction of high-energy particles with the detectors and surrounding material. In the case of observations carried out at stratospheric balloon altitudes, we also need to take into account the atmospheric radiation produced by the interaction of cosmic particles with atmospheric constituents, which will generate secondary particles and photons. In any case, the observations of cosmic X- and γ -ray sources are always hindered by intense and complex background radiation measured by the detector system. In this paper, we consider the contribution of photons, protons, electrons and neutrons to the total background of a hard X-ray coded-aperture imaging camera to be mounted first in a balloon platform and later in a satellite bus at near-equatorial low-Earth orbit (LEO).

In order to estimate the background, it is important to have an accurate knowledge of its origin. To achieve this, one needs to include detailed descriptions of the energy spectra and angular distribution

of the particle fields that surround the instrument. In this work, we describe the procedures we used in order to model the background of a hard X-ray imaging camera being developed at the National Institute for Space Research (INPE), Brazil, in the context of the *MIRAX (Monitor e Imageador de RAios X)* space astronomy mission (Braga et al. 2004; Braga 2006).

Since the instrument is supposed to fly on-board stratospheric balloons over Brazil, at an altitude of ~ 42 km and a latitude of $\sim -23^\circ$ S, and as part of a satellite experiment in a near-equatorial circular LEO, we have included both environments in our simulations. We have developed a detailed mass model of the camera and simulated all the interactions in the instrument components and materials using the *GEANT4 (GEometry ANd Tracking)* code developed at Conseil Européen pour la Recherche Nucléaire (CERN) (Agostinelli et al. 2003). The results of these simulations have allowed us to estimate with good precision the spectral response of the CdZnTe (CZT) detectors of the camera and the spatial distribution of counts over the detector plane. Also, by running simulations with different configurations and analysing the results, we were able to improve the passive shielding configuration of the camera in order to minimize the leakage of fluorescence radiation in the shield and achieve a near-uniform distribution of the background over the detector plane. This is extremely important for the quality of the reconstructed coded-mask images.

The balloon version of the instrument is called *protoMIRAX* (Braga et al. 2015) and will be launched for the first time from Cachoeira Paulista, SP, Brazil, in late 2016 for an ~ 24 -h

*E-mail: manuel.castro@inpe.br (MC); joao.braga@inpe.br (JB)

flight. The balloon flight will test several *MIRAX* subsystems in a near-space environment and will carry out imaging demonstrations of the camera by observing bright X-ray sources.

This paper is organized as follows: in Section 2, we describe the X-ray camera and present the mass model that we have developed to represent it in the simulations. In Section 3, we describe the different particle and photon spectra, as well as their angular dependence, that we have used as inputs to the simulations of the various components of the background in both the high atmospheric and space environments. In Section 4, we show the results of the simulations in terms of detector spectral response and the spatial distribution of counts over the detector plane. In Section 5, we present the results of imaging reconstruction using simulated observations of the Crab nebula and the Galactic Centre (GC) regions. Finally, in Section 8, we discuss the results obtained and present our conclusions.

2 THE X-RAY CAMERA

2.1 Description

The *protoMIRAX* balloon experiment consists of a wide-field coded-mask hard X-ray camera operating in a stratospheric balloon platform (gondola), together with auxiliary instrumentation. The experiment is described in detail by Braga et al. (2015). For *MIRAX*, it is envisaged that a space-qualified version of the same camera will operate in a satellite platform. Even though *MIRAX* is an approved concept instrument for the Brazilian space programme, the satellite bus that will carry *MIRAX* instruments are not yet fully defined.

The X-ray camera consists essentially of a coded mask, a passive shielding system surrounding the detectors, a collimator and a detector plane. The coded mask is based on a Modified Uniformly Redundant Array 13×13 basic pattern (Gottesman & Fenimore 1989), extended to 25×25 elements, placed 650 mm away from the detection plane and supported by an acrylic frame. The mask elements are made of 1.0 mm-thick lead, 1.5 mm-thick tin and 0.5 mm-thick copper, each element having an area of $20 \times 20 \text{ mm}^2$. This will provide a $\sim 20^\circ \times 20^\circ$ fully coded FoV, with an angular resolution of $\sim 1^\circ 45$ arcmin. The detection plane consists of an array of 13×13 CZT (Cadmium Zinc Telluride) crystals of dimensions $10 \text{ mm} \times 10 \text{ mm} \times 2 \text{ mm}$ with a gap of 10 mm between two consecutive crystals. The energy range is from 10 to 200 keV, although for the balloon version, we will start from 30 keV due to the high atmospheric absorption below these energies. The full detection area is thus 169 cm^2 . The collimator is a grid of crossed walls, 81 mm high, made of lead, tin and copper layers in a configuration Cu(0.2 mm)–Sn(1.0 mm)–Pb(0.5 mm)–Sn(1.0 mm)–Cu(0.2 mm). There is one detector at the bottom of each collimator cell, so that the centres of contiguous collimator blades are 20 mm apart. The collimator was designed to minimize background and define the total FoV to match exactly the fully coded FoV defined by the mask. This concept has the advantage of having zero partially coded FoV, but on the other hand, the instrument sensitivity falls off with the collimator response towards the edges of the FoV. With these parameters, as explained in Braga et al. (2015), the total, fully coded FoV of $20^\circ \times 20^\circ$ has a $7^\circ \times 7^\circ$ region of full sensitivity and $14^\circ \times 14^\circ$ full width at half-maximum (FWHM).

2.2 The mass model

In order to run the Monte Carlo simulations using *GEANT4*, we need to have photons and particles coming from random directions and energies interact with the instrument, according to pre-defined

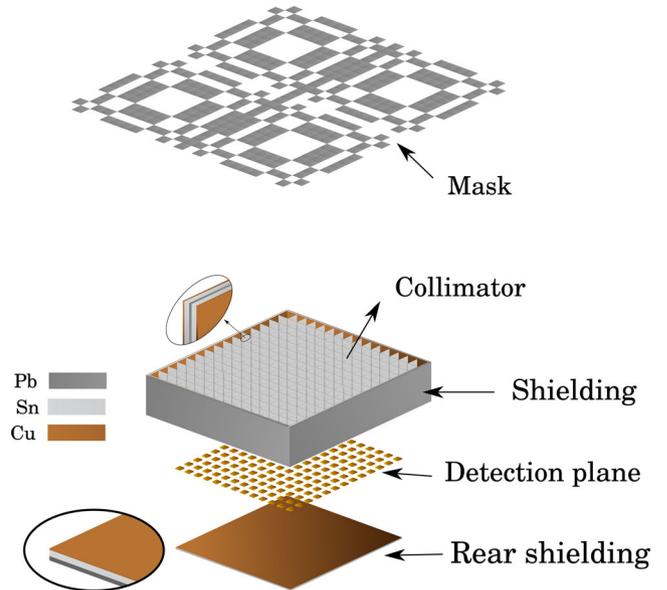


Figure 1. Mass model of the X-ray camera as mounted in *GEANT4*.

spectral and angular distributions. Therefore, we need to build a detailed description of the overall geometry of the experiment and the material constituents of each part. We also have to define the geometric shape of each single piece, their dimensions, the materials they are made of and the relative position and orientation with respect to the other parts. The code then simulates all the interactions of the primary and secondary particles with the different parts of the experiment before reaching the detectors. At the end of a simulation run, we have information on how many particles of each kind reach each of the detectors, and on the energy deposited by each particle. With this information, we can build the event distribution over the detector plane for each energy range of interest. By considering the most important radiation fields present in the environments at which the experiment will be operating, we can calculate the contribution of each component to the total background to be measured by the experiment. Furthermore, we can include hard X-ray photons from cosmic sources of interest to build simulated shadowgrams of selected source fields and then produce simulated images of astrophysical regions. In Fig. 1, we show the mass model of the hard X-ray camera as implemented in *GEANT4*.

3 ENVIRONMENTS AND FIELDS

3.1 The balloon environment

At the stratospheric environment at which the balloon flights will be carried out (altitude $\sim 42 \text{ km}$, latitude $\sim -23^\circ$), we have included in the simulations the main radiation fields: atmospheric γ -ray photons, electrons, atmospheric neutrons, and primary and secondary protons. The atmospheric γ -ray emission contribution has a dependence on the zenith angle z , so we have simulated each component separately according to the different incident angular ranges with respect to the zenith (Gehrels 1985, see Table 1): the electron spectrum in the energy range 1–10 MeV at an atmospheric depth of 5 g cm^{-2} can be modelled with a single power law (Gehrels 1985):

$$\frac{dN}{dE} = 1.4 \times 10^{-2} E^{-1.8} \quad \text{electrons cm}^{-2} \text{ s}^{-1} \text{ sr}^{-1} \text{ MeV}^{-1}. \quad (1)$$

Table 1. Atmospheric γ -ray spectrum at an optical depth of 3.5 g cm^{-2} . The power-law spectra are in units of photons $\text{cm}^{-2} \text{ s}^{-1} \text{ sr}^{-1} \text{ MeV}^{-1}$.

Spectrum	Energy range (MeV)	Angular range ($^\circ$)
$2.19 \times 10^2 E^{0.70}$	0.024–0.035	0–65
$5.16 \times 10^{-2} E^{-1.81}$	0.035–10	0–65
$0.085 E^{-1.66}$	0.1–10	65–95
$0.14 E^{-1.50}$	0.1–10	95–130
$0.047 E^{-1.45}$	0.1–10	130–180

The proton contribution has two components, primary and secondary (Dean et al. 2003), depending on the energy range and whether the protons are part of the galactic cosmic rays or result from atmospheric interactions.

The primary proton spectrum, in units of protons $\text{cm}^{-2} \text{ s}^{-1} \text{ MeV}^{-1}$, is well described by

$$\frac{dN}{dE} = \begin{cases} 1.3 \times 10^3 E^{-2} & \text{for } 10\,000 \text{ MeV} \leq E \leq 30 \text{ GeV} \\ 1.8 \times 10^6 E^{-2.7} & \text{for } 30 \text{ GeV} \leq E \leq 150 \text{ GeV} \end{cases}, \quad (2)$$

whereas for the secondary protons, the spectrum is given by

$$\frac{dN}{dE} = \begin{cases} 1.475 \times 10^{-6} E^2 & \text{for } 5 \text{ MeV} \leq E \leq 16 \text{ MeV} \\ 3.78 \times 10^{-4} E^0 & \text{for } 16 \text{ MeV} \leq E \leq 100 \text{ MeV} \\ 6.43 \times 10^{-3} E^{-0.61} & \text{for } 100 \text{ MeV} \leq E \leq 300 \text{ MeV} \\ 1.78 E^{-1.6} & \text{for } 300 \text{ MeV} \leq E \leq 6.6 \text{ GeV} \end{cases}. \quad (3)$$

The neutron background can be represented as follows (Gehrels 1985):

$$\frac{dN}{dE} = \begin{cases} 0.26 E^{-1.3} & \text{for } 0.8 \leq E \leq 10 \text{ MeV} \\ 6 \times 10^{-2} E^{-0.65} & \text{for } 10 \leq E \leq 100 \text{ MeV} \end{cases}, \quad (4)$$

where dN/dE is in units of neutrons $\text{cm}^{-2} \text{ s}^{-1} \text{ MeV}^{-1}$.

Fig. 2 shows the spectra for the different particle species that were used as input in the *protoMIRAX* background simulations.

3.2 The *MIRAX* low-Earth near-equatorial orbit environment

In order to simulate the background for the *MIRAX* observations, we have considered the main background sources at the expected satellite orbit. We exclude the satellite passages in the South Atlantic

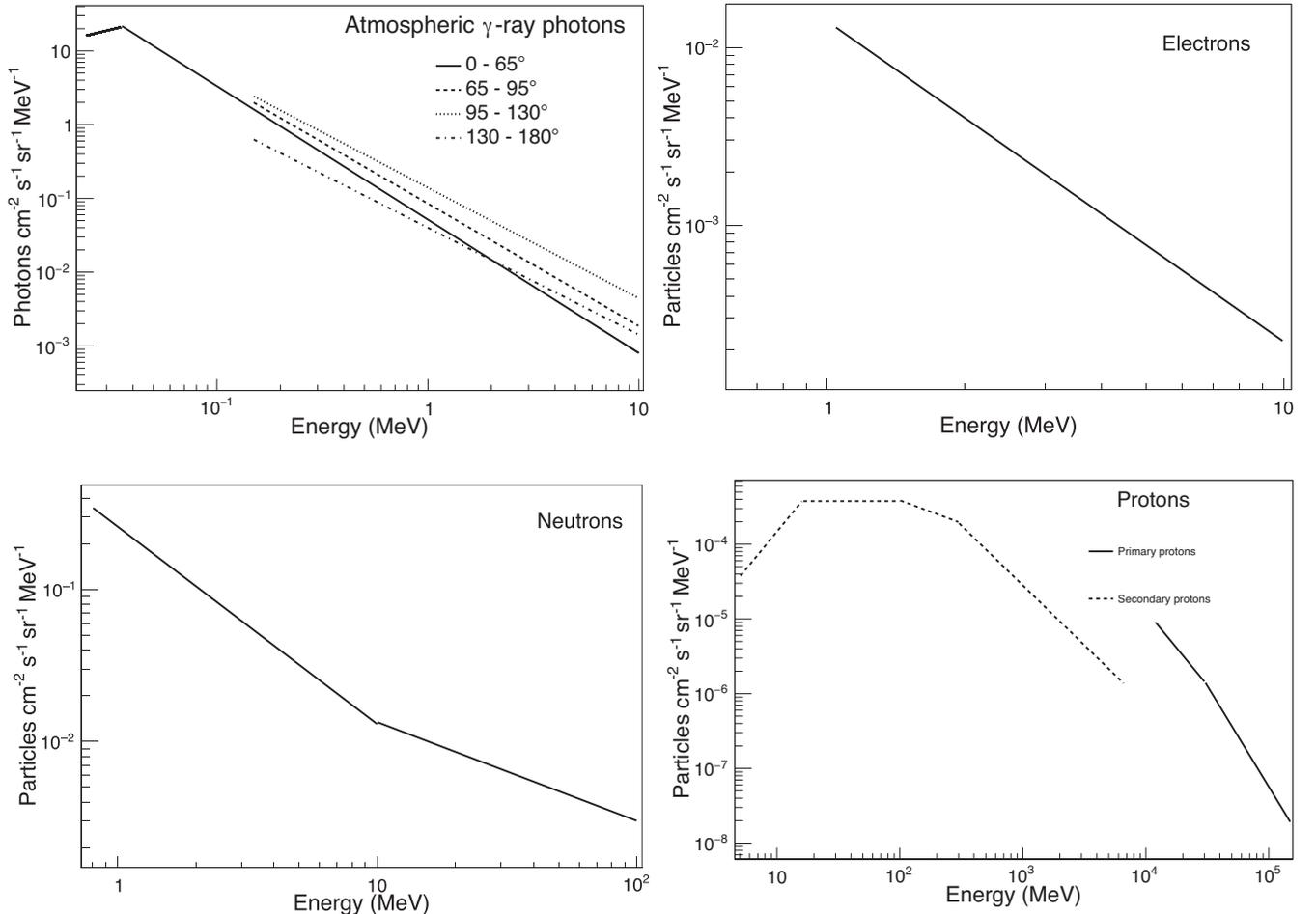


Figure 2. Background sources at balloon altitudes. Top left: γ -ray atmospheric emission. The solid line corresponds to incident photons with zenith angles between 0° and 65° ; dashed to zenith angles between 65° and 95° ; dotted between 95° and 130° , and dash-dotted to the range 130° – 180° . Top right: electron spectrum. Bottom left: neutron spectrum. Bottom right: proton spectra; solid line refers to primary protons, dashed line to secondary protons.

Anomaly (SAA), in which the particle fields are much more intense. We have considered the cosmic diffuse γ -ray radiation, albedo γ -ray photons, galactic cosmic rays (GCR), trapped protons and neutrons. The electron contribution is negligible.

The cosmic diffuse γ -ray radiation spectrum between 10 keV and 100 MeV is given by (Gruber et al. 1999):

$$\frac{dN}{dE} = \begin{cases} 7.877 E^{-1.29} e^{-E/41.43} & \text{for } E \leq 60 \text{ keV} \\ 4.32 \times 10^{-4} \left(\frac{E}{60}\right)^{-6.5} + \\ 8.4 \times 10^{-3} \left(\frac{E}{60}\right)^{-2.58} + \\ 4.8 \times 10^{-4} \left(\frac{E}{60}\right)^{-2.05} & \text{for } E \geq 60 \text{ keV} \end{cases} \quad (5)$$

To simulate the albedo γ -ray photon field, we have used the spectrum, also between 10 keV and 100 MeV, given by Ajello et al. (2008) and Sarkar et al. (2011):

$$\frac{dN}{dE} = \begin{cases} \frac{1.87 \times 10^{-2}}{\left(\frac{E}{33.7}\right)^{-5} + \left(\frac{E}{33.7}\right)^{1.72}} & \text{for } E \leq 200 \text{ keV} \\ 1.01 \times 10^{-4} \left(\frac{E}{\text{MeV}}\right)^{-1.34} & \text{for } 200 \text{ keV} \leq E \leq 20 \text{ MeV} \\ 7.29 \times 10^{-4} \left(\frac{E}{\text{MeV}}\right)^{-2} & \text{for } E \geq 20 \text{ MeV} \end{cases} \quad (6)$$

The galactic cosmic ray spectrum in the energy range 10 MeV–20 GeV is well described by Mizuno et al. (2004) and Sarkar et al. (2011):

$$\frac{dN}{dE} = F \left(\frac{E}{\text{GeV}}\right)^{-a} \exp\left(-\frac{E}{E_{1\text{cut}}}\right)^{-a+1} + A \left(\frac{E + Ze\phi}{\text{GeV}}\right)^{-b} \times \frac{(E + Mc^2)^2 - (Mc^2)^2}{(E + Mc^2 + Ze\phi)^2 - (Mc^2)^2} \times \frac{1}{1 + \left(\frac{E}{E_{2\text{cut}}}\right)^{-12}}, \quad (7)$$

where e is the electron charge, M is the particle's mass, c is the speed of light, Z is the atomic number, $F = 1.23 \times 10^{-8}$ counts $\text{cm}^{-2} \text{s}^{-1} \text{sr}^{-1} \text{keV}^{-1}$, $E_{1\text{cut}} = 5.1 \times 10^6$ keV, $a = 0.155$, $A = 2.39 \times 10^{-6}$ counts $\text{cm}^{-2} \text{s}^{-1} \text{sr}^{-1} \text{keV}^{-1}$, $b = 2.83$, $\phi = 6.5 \times 10^5$ kV and $E_{2\text{cut}} = 1.13 \times 10^7$ keV.

To simulate the trapped proton contribution from 100 MeV to 6 GeV, we have used the following expression (Mizuno et al. 2004; Sarkar et al. 2011)

$$\frac{dN}{dE} = F \left(\frac{E}{\text{GeV}}\right)^{-a} \exp\left(-\frac{E}{E_{1\text{cut}}}\right)^{-a+1}, \quad (8)$$

with the constant values the same as above.

Finally, the neutron spectrum between 10 keV and 1 GeV at these altitudes was modelled according to (Armstrong, Chandler & Barish 1973):

$$\frac{dN}{dE} = \begin{cases} 9.98 \times 10^{-8} \left(\frac{E}{\text{GeV}}\right)^{-0.5} & \text{for } 10 \text{ keV} \leq E \leq 1 \text{ MeV} \\ 3.16 \times 10^{-9} \left(\frac{E}{\text{GeV}}\right)^{-1.0} & \text{for } 1 \text{ MeV} \leq E \leq 100 \text{ MeV} \\ 3.16 \times 10^{-10} \left(\frac{E}{\text{GeV}}\right)^{-2.0} & \text{for } 100 \text{ MeV} \leq E \leq 100 \text{ GeV} \end{cases} \quad (9)$$

For all the equations above (5, 6, 7, 8 and 9), dN/dE is given in particles $\text{cm}^{-2} \text{s}^{-1} \text{sr}^{-1} \text{keV}^{-1}$.

Fig. 3 shows the spectra for the different kinds of particles in the respective energy ranges.

4 THE SIMULATED BACKGROUND

Using the mass model and the radiation spectra described above as inputs to GEANT4, we performed simulations to study in detail the

instrumental response of the X-ray camera under the two environments described in the previous section.

For each radiation field, we have used closely spaced data points from the original spectrum to run separate simulations for each component. GEANT4 runs internal interpolation tasks to generate random events following the spectral behaviour of the original spectrum. In order to comply with the GEANT4 specific input procedures, we have built a model in which the X-ray camera is placed at the centre of a virtual emission sphere. The emissions from the spherical surface are then randomly created by the programme both in terms of energy and directions, always pointing to the internal volume of the sphere. Primary particles coming from space will reach the experiment from any direction between $z = 0$ and $z = \pi/2$, where z is the zenith angle, while secondary particles, generated by the interaction of primary particles with the atmosphere, can reach the instruments from all directions. For *protoMIRAX*, the atmospheric gamma radiation spectra are generated in spherical sectors bounded by the incident angle ranges described above, while the primary protons come from the upper hemisphere and the neutrons, secondary protons and electrons come equally from all directions. In this case, the radius of the virtual sphere (300 cm) was very large large compared to the instrument dimensions in order to insure that the randomly generated events in the internal surface of the sphere would follow closely the angular distribution of the radiation field without introducing significant effects due to the local geometry. For the *MIRAX* environment, all the events hitting the detector are coming from a hemisphere, either centred on the zenith (diffuse γ -rays and cosmic rays) or the nadir (the other three fields). The radius of the sphere used in this case (90 cm) did not have to be too large since effects introduced by the local geometry are not significant.

From the GEANT4 output files, we can determine the deposited energy in each of the detectors by each kind of interaction chain in the instrument. The output of the simulations allows us to build a spectrum and a spatial distribution of events over the detector plane for each of the radiation fields that interact with the instrument and for each of the 169 CZT detectors. Those components can then be added up to represent the total background to be measured by *protoMIRAX* and *MIRAX*. This separation of the background contributions is extremely useful for the design of the structure and geometry of the instrument's passive shielding. By using the best shielding configuration, we can homogenize the spatial distribution of counts over the detector plane and lower the overall background. This is essential to improve the instrument sensitivity.

4.1 Normalization

In the simulations using GEANT4, the user specifies the number N of virtual particles that will be created, either coming from random directions or from user-specified directions. In the case of a point source at infinity, like for instance a stellar X-ray source, the photons come in parallel lines at random positions inside a cylinder that encompasses the whole instrument. The lines are all parallel to the cylinder's axis. For radiation fields that have angle dependence, the user has to specify a spherical surface from which the particles will come, as explained in the previous section. GEANT4 creates virtual particles at random positions in the surface of the sphere and with random directions, provided they all point to the inside volume of the sphere. The programme then creates the secondary, tertiary etc. particles that are produced by all interactions in the instrument mass model and tracks all their directions. Finally, it computes the energy losses on the detectors produced by the interactions of all particles that hit each specific detector's volume.

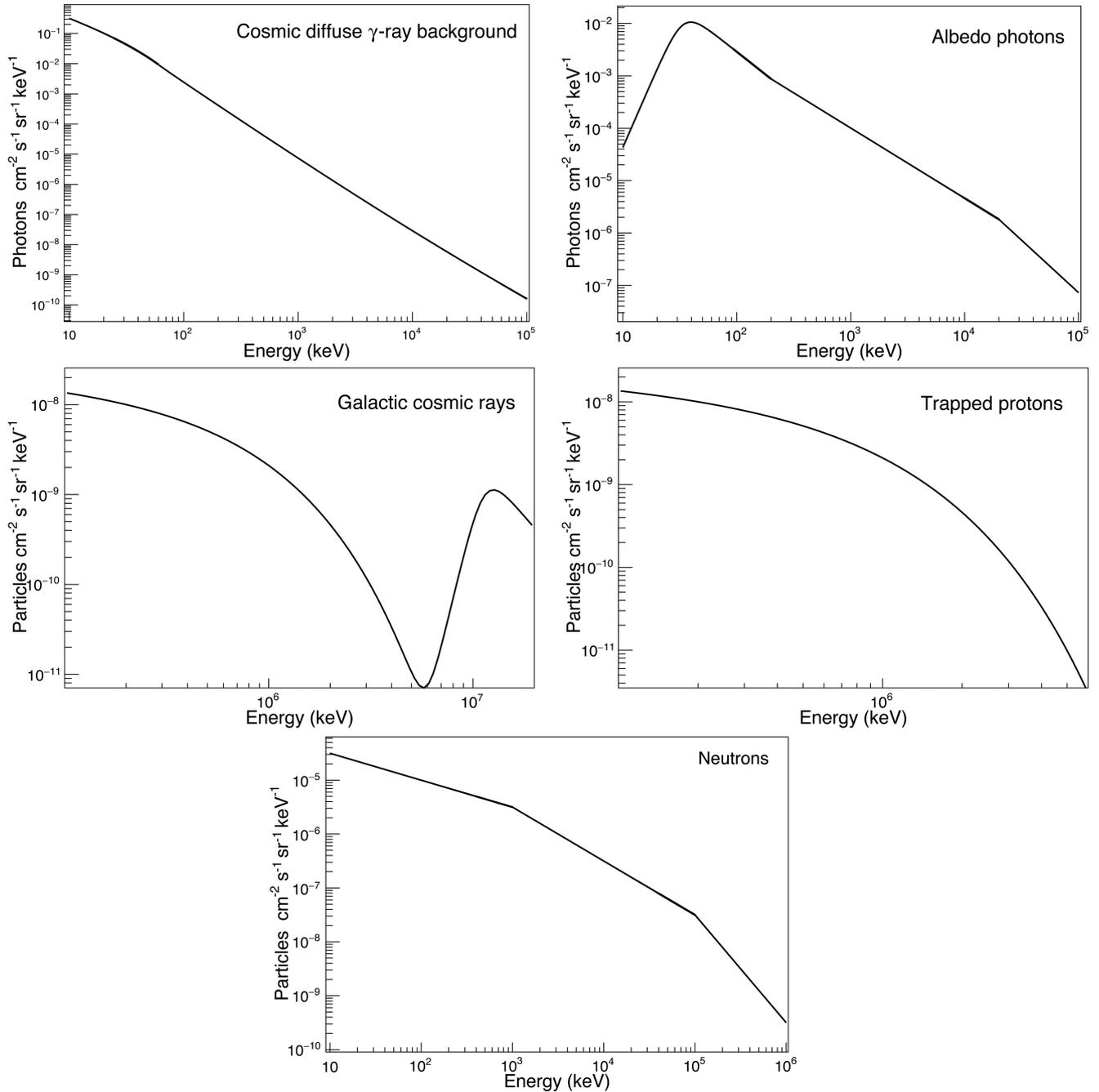


Figure 3. Spectra of each radiation field present at the *MIRAX* orbit (outside the SAA), used as input in the *GEANT4* simulations of the background.

One important question that arises is how to relate the integration time of the simulated observation to the value of N . In our simulations, we have used the following procedure. The geometrical scheme is shown in Fig. 4. We first note that the number of particles dN passing through the area dA within a solid angle $d\Omega$ during a time dT in the energy range dE is $dN = I_N dA dT d\Omega dE$, where I_N is the specific intensity of the radiation field in units of particles $\text{cm}^{-2} \text{s}^{-1} \text{sr}^{-1} \text{keV}^{-1}$. We now define an element of area dA in an arbitrary direction \hat{n} . The net differential flux from the solid angle $d\Omega$ is given by $dF_N = I_N \cos \theta d\Omega$. The net flux is then obtained integrating dF_N over all the solid angles,

$$F_N = \int I_N \cos \theta d\Omega. \quad (10)$$

Now, the total number of particles N created in the sphere (or a spherical sector) to represent this flux will be

$$N = \int I_N \cos \theta d\Omega dA dt dE, \quad (11)$$

with $d\Omega = d\phi \sin \theta d\theta$, $dA = R^2 \sin \alpha d\alpha d\Phi$ (see Fig. 4), $0 \leq \theta \leq \pi/2$ and $0 \leq \phi \leq 2\pi$. R is the radius of the sphere and the range of α will depend on the specific spherical sector related to the incident angles of the radiation field.

We then have

$$N = \pi A T \int I_N dE, \quad (12)$$

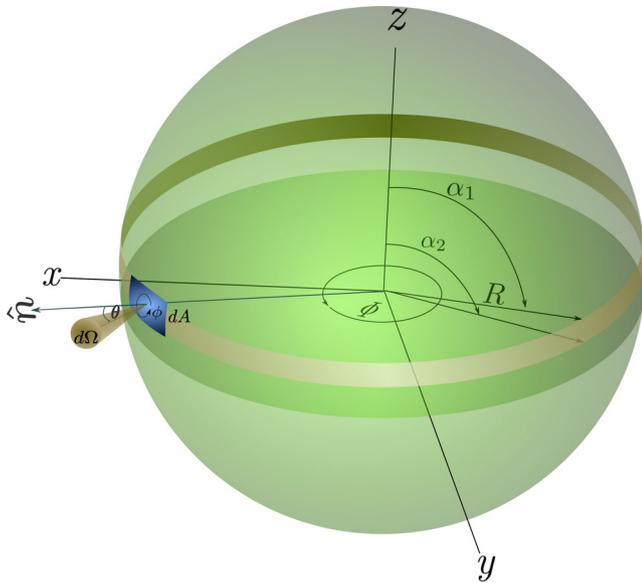


Figure 4. Geometry used for the normalization of GEANT4 simulations.

where A is the geometrical area of the spherical sector and T is the total integration time. Finally, the integration time is obtained by

$$T = \frac{N}{\pi A \int I_N dE}, \quad (13)$$

where $I_N = dN/dE$ is directly related to the radiation field spectra of Sections 3.1 and 3.2. For the *protoMIRAX* experiment, we have considered an 8-h integration time in order to have statistics compatible with what is achieved in time-scales of balloon flights. For the satellite version, each radiation spectrum (see Fig. 2) was considered separately so as to have good enough statistics for a good estimation of all background components.

4.2 Background spectra

In Fig. 5, we show the simulated spectra produced by the several radiation components at balloon altitudes. One can see that the contribution of X-rays from atmospheric ‘fast’ neutron interactions in the instrument is very important below ~ 50 keV, whereas the contributions of photons, protons and neutrons are very similar between ~ 100 and ~ 150 keV. The atmospheric diffuse gamma

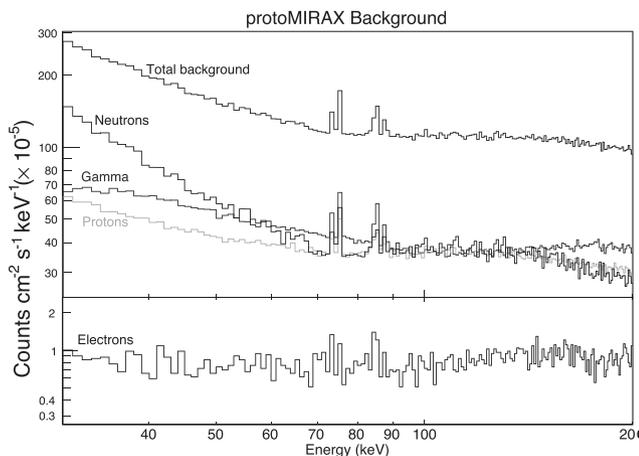


Figure 5. Total background spectrum measured by *protoMIRAX*.

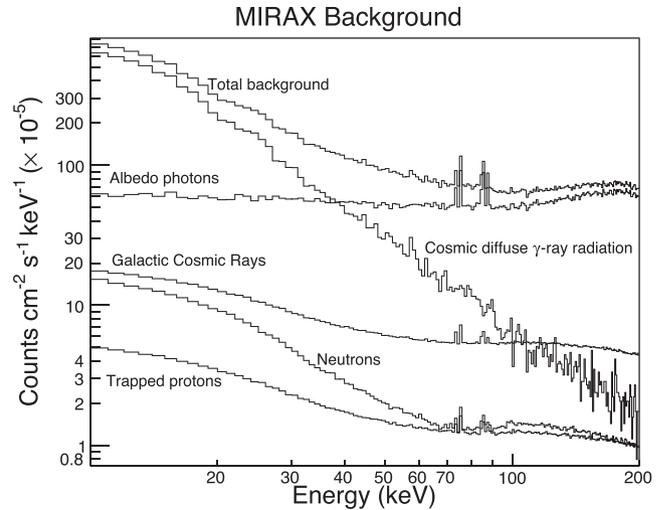


Figure 6. Total background spectrum measured by *MIRAX*. The contribution of each individual component is also shown.

radiation is the strongest component above ~ 150 keV. The contribution from electrons, about two orders of magnitude weaker, is clearly negligible. The fluorescence lines of lead are clearly visible at 73, 74, 84 and 85 keV (The National Physical Laboratory 2008). We have changed the thicknesses of the passive shielding materials (Pb–Sn–Cu) several times until we reached the best solution for the instrument in terms of leakage of the fluorescence lines without adding too much weight.

In Fig. 6, we show the camera’s background spectrum at the orbit of *MIRAX*. The cosmic diffuse radiation is the dominant component up to ~ 40 keV. The interactions of albedo photons from the atmosphere below contribute very strongly and becomes a very important component all the way to 200 keV. Galactic cosmic rays (essentially protons) are also important, albeit about one order of magnitude weaker. It is important to mention that currently we are optimising the X-ray camera for *protoMIRAX*. After the balloon flights, we will make the necessary changes in the configuration so as to optimize the camera for operation in orbit. The contribution of albedo photons will certainly be greatly reduced by a different shielding system and by spacecraft materials as well. The simulations reported here will be of utmost importance to the final design of the camera to operate in space.

4.3 Background spatial distribution over the detector plane

A very important result of our simulations is the spatial distribution of background counts across the detector plane. In the case of *protoMIRAX*, we have taken into account the fact that the γ -ray flux at the upper atmosphere is highly anisotropic. In order to obtain an accurate description of what happens at the detector level for each of the angular components, we produced count maps for each component (first four plots in Fig. 7). In the 0° – 65° range, one can see the effects of the collimator response: since the flux from 0° to 3° is higher than at larger angles due to the collimator response (see Braga et al. 2015), the shadow of the central bar on the mask pattern dominates so that we have fewer counts on the detectors underneath it. We have confirmed this hypothesis by rotating the mask 90° . For the 65° – 95° range, we see a deep non-uniform distribution due to the fact that the detectors near the corners and edges of the detector plane are less protected by the collimator blades and shielding materials. γ -rays coming from below the instrument have a nearly

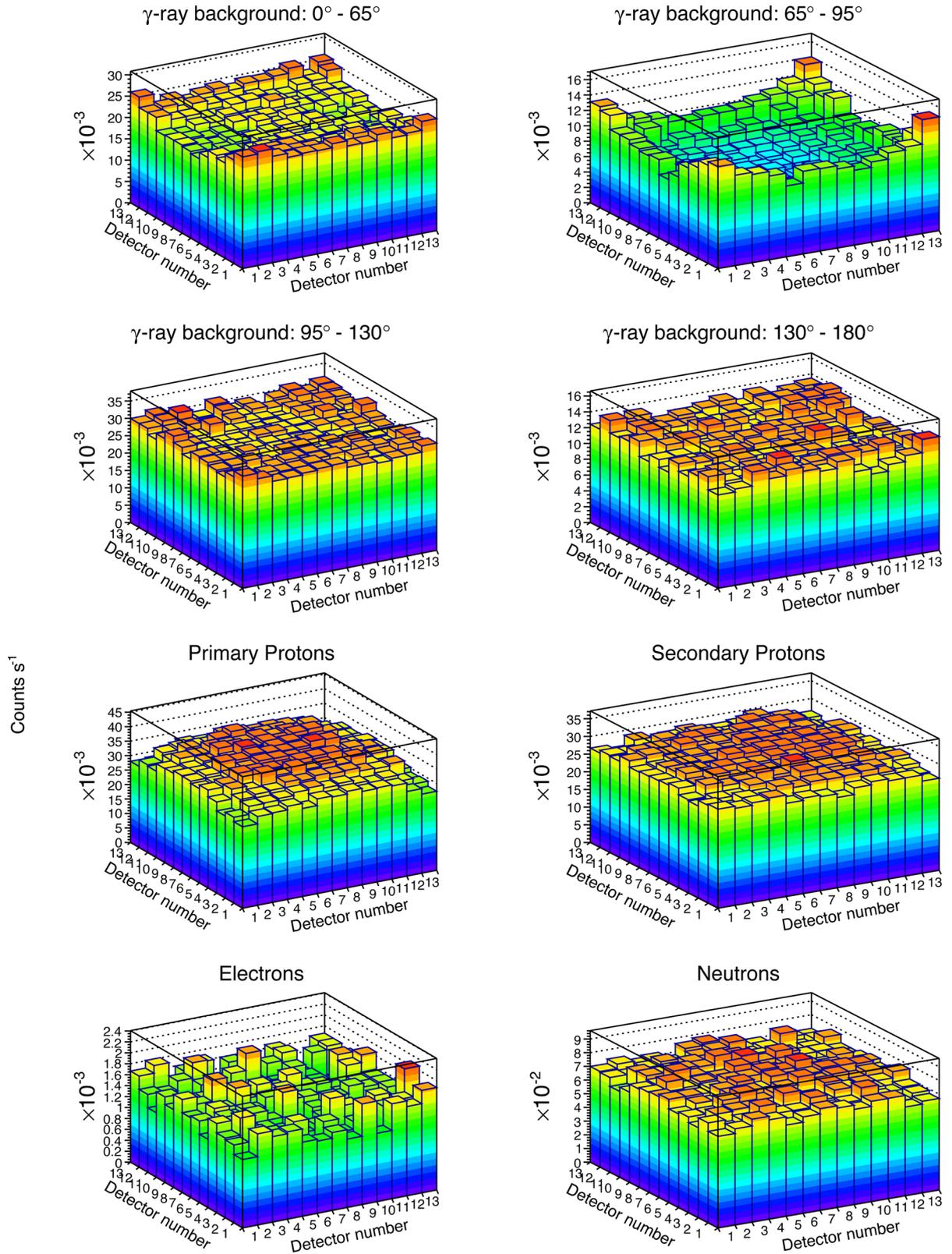


Figure 7. Count rate distribution for each background component, over the full detection plane for *protoMIRAX*. The simulated integration time is 8 h for all distributions, and the energy range is 30–200 keV. The typical 1σ error bars (left-right, top-bottom) for the detectors are 8.8×10^{-4} , 5.2×10^{-4} , 1.0×10^{-3} , 6.7×10^{-4} , 1.0×10^{-3} , 1.0×10^{-3} , 2.2×10^{-4} , and 1.6×10^{-3} counts s^{-1} .

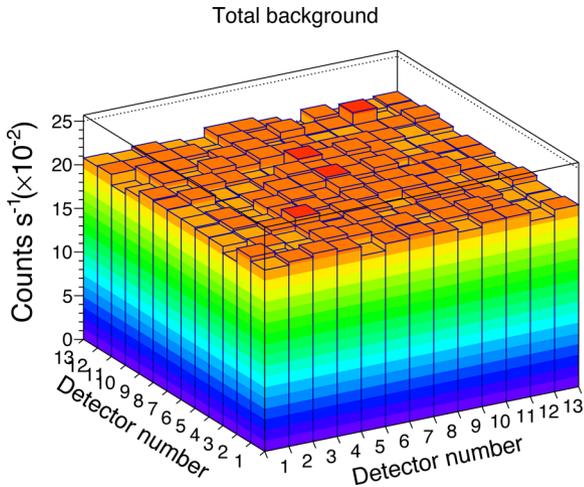


Figure 8. Count rate distribution of the total background over the detection plane in the energy range 30–200 keV. The integration time is 8 h at balloon altitudes. The typical 1σ error bar for the detectors is 2.8×10^{-3} counts s^{-1} .

homogeneous distribution, as expected. The distribution caused by primary protons shows less counts towards the edges and corners. This is also expected, since the probability for production of secondary radiation increases with the number of collimator blades the protons interact with. This is less pronounced with secondary protons, since they come from all sides. The atmospheric neutrons and electrons produce a more uniform distribution, within statistics. If we add all components of the background (Fig. 8), we get a very uniform distribution. This was also a consequence of the particular configuration we have implemented for the collimator and shielding materials, as reported by Braga et al. (2015). The total (30–200 keV) background count rate for *protoMIRAX* is estimated to be 35.5 counts s^{-1} . It is noteworthy that preliminary results reported by Penacchioni et al. (2015) have provided a count rate of ~ 50 counts s^{-1} . The improved mass model we built based on preliminary simulation results have led to this 29 per cent decrease in the background level.

For *MIRAX*, the distributions (Fig. 9) show similar structures. In this case, the predominance of the diffuse γ radiation up to 40 keV makes the total background distribution less homogeneous. The count map due to albedo photons coming from below, which dominates above 40 keV, is uniform within statistics. In this case, the count rate is 35.3 counts s^{-1} in the total energy range 10–200 keV. The similar result with respect to *protoMIRAX* is a coincidence due to the fact that the background is lower but the energy band is wider. In the 30–200 keV range, the same band of *protoMIRAX*, the total count rate for *MIRAX* is 22.4 counts s^{-1} .

5 IMAGING RECONSTRUCTION

In order to study the effects of the background in the imaging reconstruction process, we have considered the presence in the FoV of point sources than can be observed by *protoMIRAX* and *MIRAX*. The sources are the Crab nebula and two well-known Galactic black holes, 1E 1740.7–2942 and GRS 1758–258, near the Galactic Centre. The Crab’s spectrum can be represented by a broken power-law model with $\Gamma_1 = 2.105$, $E_{\text{break}} = 100$ keV, $\Gamma_2 = 2.22$ and a flux at 1 keV of 10.2 photons $\text{cm}^{-2} \text{s}^{-1} \text{keV}^{-1}$ (Jourdain et al. 2008).

As 1E 1740.7–2942 and GRS 1758–258 spend most of their time in the low/hard state (Remillard & McClintock 2006), their

emission in X/ γ -rays can be represented as a combination of a power law and a disc emission (if necessary). The model parameters for each source are shown in Table 2. The three spectra are shown in Fig. 10.

5.1 Imaging reconstruction for *protoMIRAX*

Since the *protoMIRAX* experiment will fly in the atmosphere at an altitude of ~ 42 km, the fluxes from the astrophysical sources will suffer atmospheric absorption according to:

$$F(E) = F_0(E) e^{-\frac{\mu}{\rho}(E) x \sec(z)}, \quad (14)$$

where F_0 is the flux at the top of the atmosphere, μ/ρ is the absorption coefficient of the atmosphere (Hubbell & Seltzer 2004) in units of $\text{cm}^2 \text{g}^{-1}$, z is the zenith angle and x is the expected atmospheric depth (2.7 g cm^{-2}) at an altitude of ~ 42 km. The zenith angle is a function of the geographical latitude φ where the balloon will fly, the right ascension α and declination δ of the observed source, and the local sidereal time t . The zenith angle is given by

$$\cos z(t) = \sin \varphi \sin \delta + \cos \varphi \cos \delta \cos(t - \alpha). \quad (15)$$

In this way, it is possible to compute the number of photons N that will hit the detector coming from a specific point source, attenuated by the atmosphere. The number N of photons passing through an area S , in a time interval T and in the energy range $E_{\text{min}}-E_{\text{max}}$ is given by

$$N = \int_S \int_{E_{\text{min}}}^{E_{\text{max}}} \int_0^T F_0(E) e^{-\frac{\mu}{\rho}(E) x \sec(z(t))} dt dE dS. \quad (16)$$

We have simulated the Crab radiation by integrating equation (16) in the energy range 30–200 keV for a meridian passage of the source at a latitude of -23° . We considered a 4-h interval centred at the meridian crossing. The Crab’s reconstructed image is shown in Fig. 11. These high signal-to-noise ratio (~ 80) images of the Crab will be used to calibrate the imaging system and demonstrate the performance of the camera in a near-space environment.

5.2 Imaging reconstruction for *MIRAX*

In the case of *MIRAX*, we have considered the three sources described previously, but in a space environment (no atmospheric absorption). The reconstructed images are shown in Figs 12 (for the Crab) and 13 (for 1E 1740.7–2942 and GRS 1758–258). We can see a great improvement in the Crab image (4 h of integration in both cases), due to the lower background in space and mainly due to the fact that the Crab is observed without atmospheric absorption by *MIRAX*. Nevertheless, it is noteworthy that the spatial non-uniformity of the background across the detector plane, mostly due to the fact that the cosmic diffuse γ -ray component is not evenly distributed (see first panel of Fig. 9), produces a ~ 40 per cent loss in the signal to noise. This was calculated by comparing our Crab image with another in which the total background is replaced by a flat count distribution (with statistical fluctuations) of the same level. The Signal-to-Noise Ratio (SNR) that we obtain in the latter is 709, whereas in the Crab image of Fig. 12, the SNR is 428.

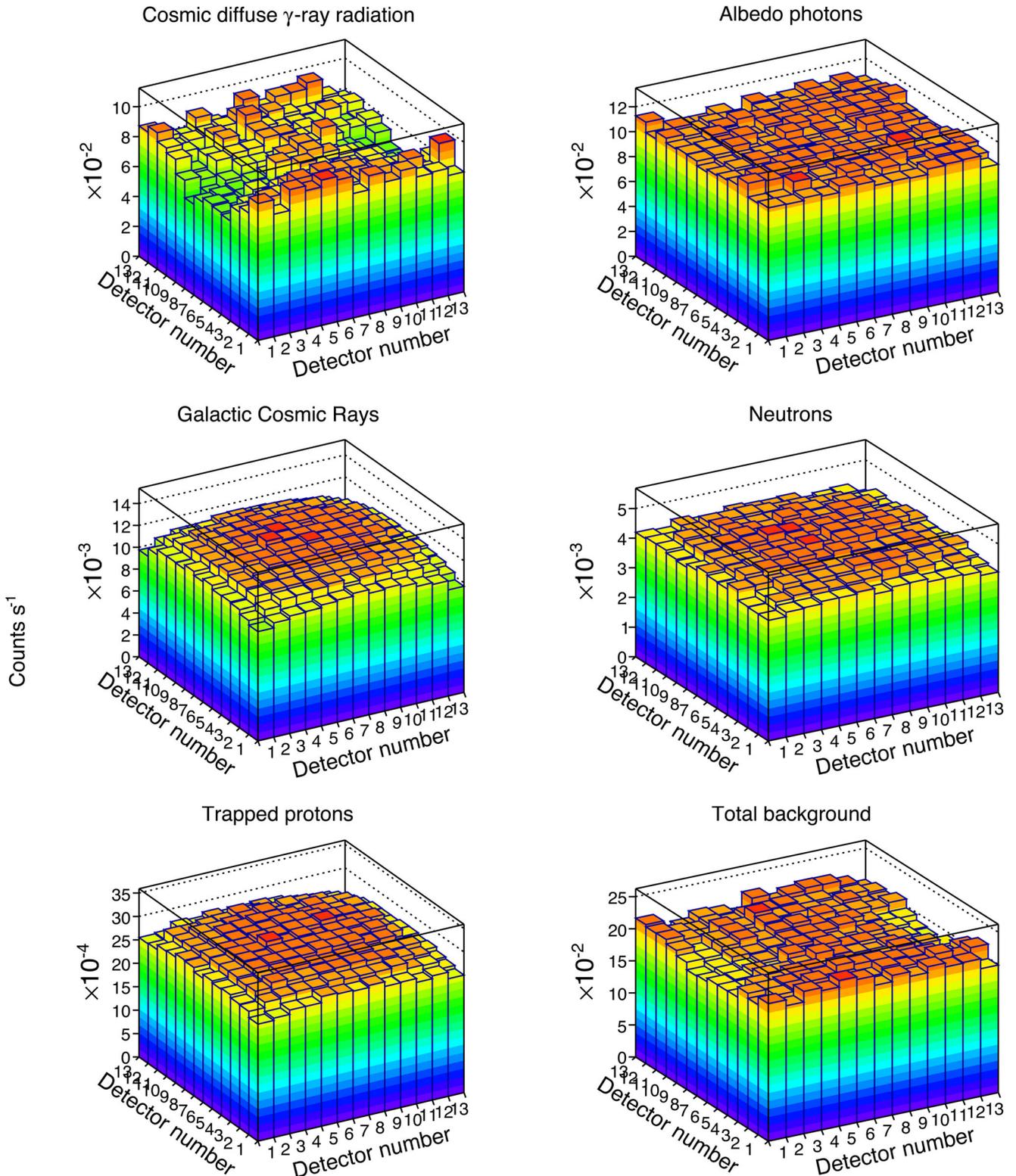


Figure 9. Count rate distributions for each background component over the detection plane for *MIRAX*. The energy range is 10–200 keV. The typical 1σ error bars, in counts s^{-1} , are: diffuse γ radiation: 4.7×10^{-3} ; albedo: 3.0×10^{-3} ; GCR: 1.2×10^{-4} ; neutrons: 8.9×10^{-5} ; trapped protons: 3.7×10^{-5} ; total background: 5.6×10^{-3} .

In the GC image, we have put 1E 1740.7–2942 in the centre of the FoV. The detected flux of GRS 1758–258 is lowered by the collimator response. It is important to note that the purpose of these images is to carry out a preliminary demonstration of the

imaging system for the X-ray camera in space, based on the simulations reported here. During the development of *MIRAX*, we are pursuing alternatives in terms of detector technology that will eventually allow us to have a detector plane with much larger area and

Table 2. Spectra used for simulating the emission of 1E 1740.7–2942 and GRS 1758–258. The values are compatible with the sources in the low/hard state.

		diskbb + cutoffpl	
		1E 1740.7–2942 ^a	GRS 1758–258 ^b
diskbb	T_{in} (keV)	0.15	0.67
	Γ	1.52	1.59
cutoffpl	E_{cut} (keV)	87.1	136

Notes. ^aCastro et al. (2014).

^bPottschmidt et al. (2006).

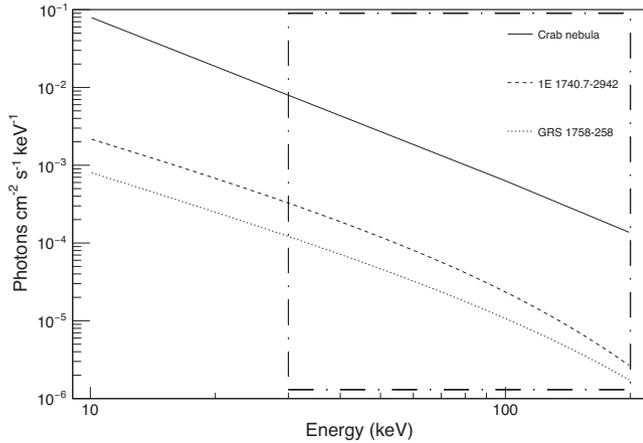


Figure 10. Spectra used to simulate the emission from Crab, 1E 1740.7–2942 and GRS 1758–258. The dashed box represents the energy range used to simulate the imaging reconstruction for *protoMIRAX* (considering attenuation by the atmosphere). For the *MIRAX* imaging reconstruction, we have considered the full energy range (10–200 keV).

higher spatial resolution. This will certainly greatly improve the performance of the camera in terms of sensitivity and angular resolution, and make it a competitive instrument capable of achieving the project’s scientific objectives described by Braga (2006).

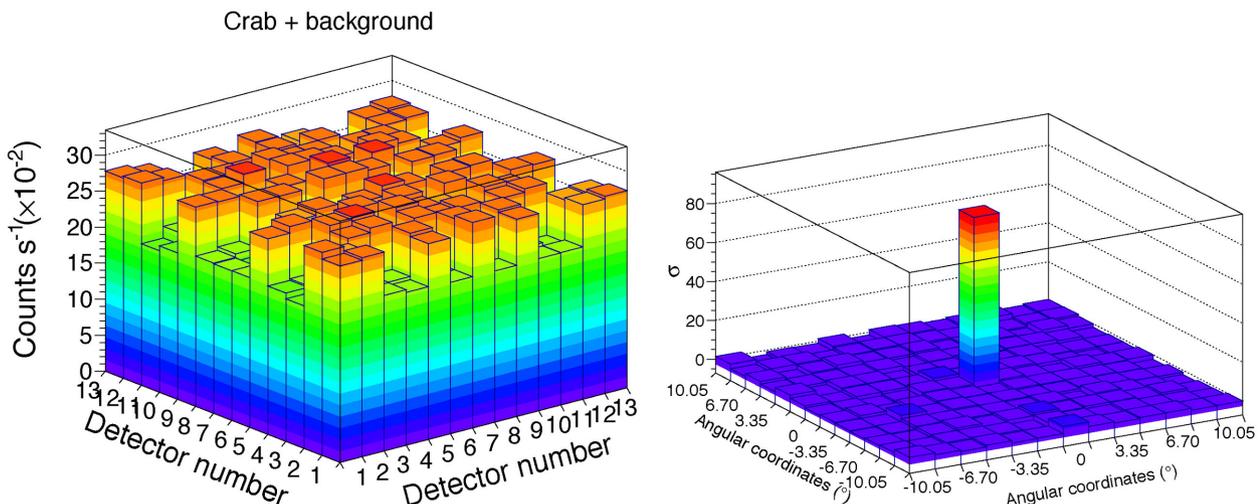


Figure 11. Crab-simulated image as observed by *protoMIRAX* during 4 h around a meridian passage at a latitude of -23° . Left: count map over the detector plane, including background. Right: reconstructed image in units of σ .

6 SENSITIVITY

In Braga et al. (2015), we have shown a sensitivity curve for *protoMIRAX*. Here we present an improved version of it based on the new background spectral distribution that we report in this work (Section 4.2). In particular, the refined mass model for the experiment, including different thicknesses for shielding materials, allowed us to decrease significantly the overall background and the intensity of the lead fluorescence lines. This has lowered the minimum observable flux, especially from 70 to 90 keV, improving the sensitivity. In this work, we have also produced a sensitivity curve for the space version of the camera (*MIRAX* mission). The sensitivity values are calculated as the minimum detectable fluxes (at a given statistical significance) for energy bins centred at each energy E :

$$F_{\text{min}} = \frac{2N_\sigma}{\epsilon(E)} \sqrt{\frac{B(E)}{A_{\text{det}} T \Delta E}} \text{ photons cm}^{-2} \text{ s}^{-1} \text{ keV}^{-1}, \quad (17)$$

where N_σ is the statistical significance in σ units (or signal-to-noise ratio, in this case), $\epsilon(E)$ is the detector efficiency at energy E (see Braga et al. 2015), $B(E)$ is the background level in counts $\text{cm}^{-2} \text{ s}^{-1} \text{ keV}^{-1}$ at energy E , A_{det} is the geometrical area of the detector plane in cm^2 , T is the integration time in seconds, and ΔE is the energy band in keV. The factor of 2 in the above expression appears due to the fact that, for a coded aperture telescope with an open mask fraction of 0.5, $A_{\text{eff}} = \epsilon A_{\text{det}}/2$, where ϵ is the detector efficiency.

The sensitivity plots are shown in Fig. 14.

7 DISCUSSION

The simulations shown in this paper have provided very important information and constraints for the optimal design of the hard X-ray camera we are planning to use as part of the *MIRAX* space programme at INPE. The configuration we are currently developing, with a collimator and a coded mask, has shown to yield interesting results essentially due to two important reasons: (1) by using a collimator, we can make the full FoV of the instrument to be coded by the mask, avoiding the well-known undesirable effects of the ubiquitous partially coded FoVs in coded aperture imaging; and (2) the mounting of the camera become simpler and lighter, since we

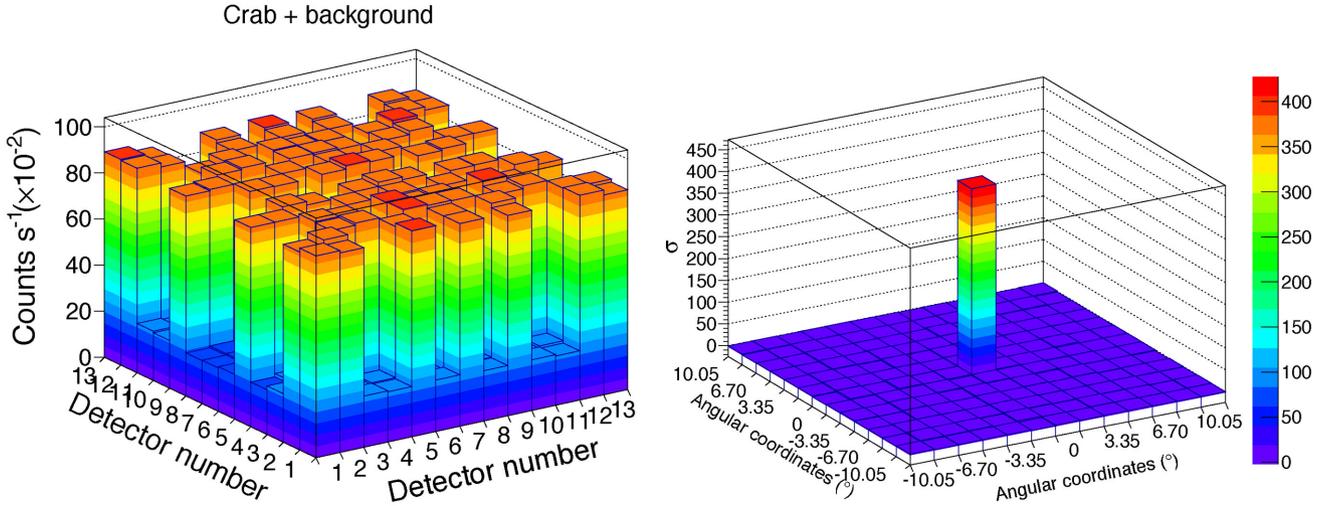


Figure 12. Imaging reconstruction of Crab as observed by *MIRAX*. Left-hand panel: distribution of counts over the detection plane considering both source and background. Right-hand panel: reconstructed image for the Crab during a 4-h observation.

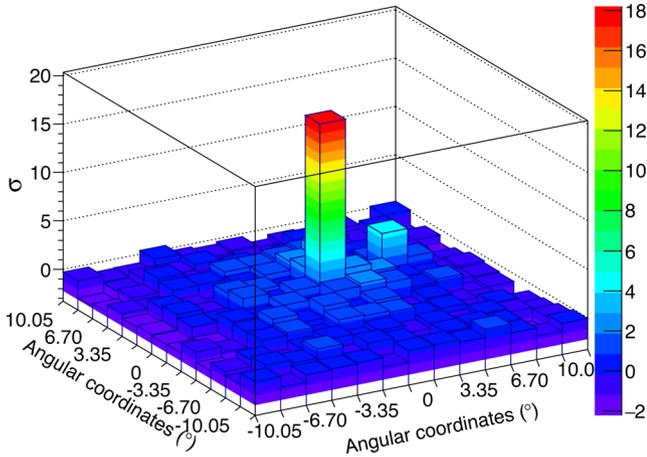


Figure 13. GC reconstructed image with 1E 1740.7–2942 (in the centre) and GRS 1758–258 in the field, for an observation of 8 h.

avoid having to include large shielding flat panels on the sides of camera, joining the detector plane with the mask; a simple slim structure made of a light material can support the mask. In our case, the weight of the collimator is about half the weight of the shielding panels we would have to build for providing roughly the same shielding effectiveness. If we build cameras with larger distances between masks and detector to achieve better angular resolution, this becomes increasingly more important. However, the use of a collimator also brings undesirable properties since the instrument's sensitivity will decrease towards the edges of the FoV according to the collimator response. In our case, the angular FWHM response provided by the collimator is 14° and imaging beyond that will be severely limited. Another problem is that the collimator produces a distribution of counts in the detector plane that is spatially non-uniform and depend on the incident directions of the components that are responsible for the detector background. We have shown here that this effect is more significant when the camera is placed in the low-orbit environment due to the intense cosmic X-ray diffuse radiation (CDR).

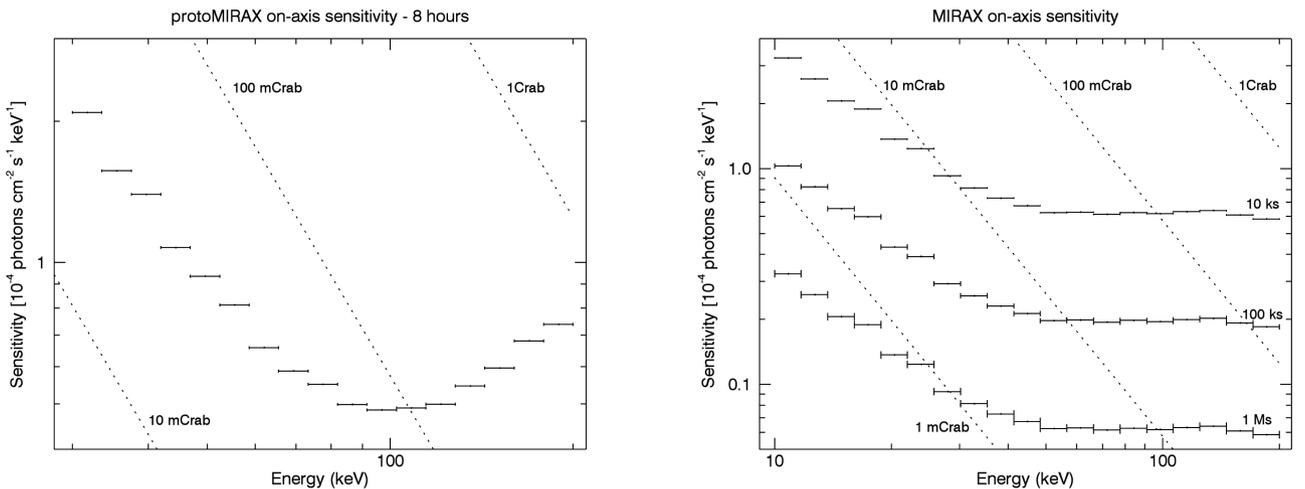


Figure 14. Left-hand panel: on-axis sensitivity curve for *protoMIRAX* for an integration of 8 h at an atmospheric depth of $2.7 g cm^{-2}$ and a latitude of -23° over Brazil. Right-hand panel: on-axis sensitivity curve for *MIRAX* for three different integration times at near-equatorial LEO. In both cases, the horizontal bars represent the minimum detectable flux at a statistical significance of 3σ .

The optimum design of a coded aperture imager will depend on the scientific objectives of the experiment. If a survey instrument is preferred, the use of collimators will probably not be the best option due to the limitations on the sensitivity over the FoV. If, on the other hand, the objective is to obtain sharp images of selected fields with no interference from adjacent fields with bright sources, a collimator could be the best option.

The simulations presented here have shown that, on hard X-ray balloon experiments, the background induced by neutrons is very important below 50 keV. Above this energy, the background produced by γ -rays and protons are equally important, and electrons make up a negligible contribution. At LEO, the CDR is by far the most important contribution to the background up to ~ 40 keV. Above that, the contribution from albedo radiation becomes dominant. The analysis of these relative contributions is very important to the design of shielding structures and materials for X-ray experiments.

8 CONCLUSIONS

The design and development of space astronomy X- and γ -ray experiments require a reliable estimation of the background levels against which the sources of interest will be observed. Particularly, in coded aperture experiments, a large detector area is fundamental for achieving competitive sensitivities. In those cases, the backgrounds are usually very intense and show inhomogeneities across the detector plane due to geometrical factors.

In this paper, we show detailed Monte Carlo simulations of the background and imaging observations of a hard X-ray imaging camera that is being developed in the scope of the *MIRAX* space mission. This instrument is a prototype that is going to be tested in stratospheric balloon flights. The *MIRAX* mission will play a very important role in the study of hard X-ray sources and transient phenomena, since it will constantly monitor a large area around the central Galactic plane region.

We present separate spectra and count distribution across the detector plane for each relevant radiation field impinging in the instrument and discuss their contribution. We also take into account the angular distribution of the different incident photons and particles. We provide detailed background characteristics and levels for two different environments in which the camera is supposed to operate: stratospheric balloon altitudes and near-equatorial LEO.

The simulations reported here allowed us to define with high accuracy the best configuration for the shielding structure and coded-mask materials of the instrument. In order to minimize the intense lead fluorescence lines in the range of ~ 70 – 80 keV that we have seen in the initial simulations, we were able to obtain an optimized configuration of materials for both the shielding and the mask. Also, the initial concentration of counts towards the edges and corners of the detector plane was completely overcome by a different geometry of both the external collimator blades and the placement of side shielding materials. The simulations have provided us with the precise count rates expected when we operate the instrument at balloon altitudes and LEO, and enabled the calculation of detailed sensitivities as a function of energy.

To the best of our knowledge, this is the first time a detailed description of `GEANT4` simulations for space astronomy instruments is reported. We hope that this will be of great value for researchers and instrument developers in this field.

ACKNOWLEDGEMENTS

We thank FINEP for financial support under Convênio 01.10.0233.00. We also thank CNPq and FAPESP, Brazil, for support under INCT Estudos do Espaço. AP acknowledges the support by the international Cooperation Programme CAPES-ICRANET financed by CAPES – Brazilian Federal Agency for Support and Evaluation of Graduate Education, within the Ministry of Education of Brazil. MC acknowledges the support of INPE/CNPq through the PCI programme and RS acknowledges the support by CAPES, Brazil.

REFERENCES

- Agostinelli S. et al., 2003, Nuclear Instruments and Methods in Physics Research Section A: Accelerators, Spectrometers, Detectors and Associated Equipment, 506, 250
- Ajello M. et al., 2008, ApJ, 689, 666
- Armstrong T. W., Chandler K. C., Barish J., 1973, J. Geophys. Res., 78, 2715
- Braga J., 2006, in D’Amico F., Braga J., Rothschild R. E., eds, AIP Conf. Ser. Vol. 840, The Transient Milky Way: A Perspective for *MIRAX*. Am. Inst. Phys., New York, p. 3
- Braga J. et al., 2004, Adv. Space Res., 34, 2657
- Braga J. et al., 2015, A&A, 580, A108
- Castro M., D’Amico F., Braga J., Maiolino T., Pottschmidt K., Wilms J., 2014, A&A, 569, A82
- Dean A. J., Bird A. J., Diallo N., Ferguson C., Lockley J. J., Shaw S. E., Westmore M. J., Willis D. R., 2003, Space Sci. Rev., 105, 285
- Gehrels N., 1985, Nucl. Instrum. Methods Phys. Res. A, 239, 324
- Gottesman S. R., Fenimore E. E., 1989, Appl. Opt., 28, 4344
- Gruber D. E., Matteson J. L., Peterson L. E., Jung G. V., 1999, ApJ, 520, 124
- Hubbell J. H., Seltzer S. M., 2004, Tables of X-Ray Mass Attenuation Coefficients and Mass Energy-Absorption Coefficients (version 1.4). Available at: <http://physics.nist.gov/xaamdi>
- Jourdain E., Götz D., Westergaard N. J., Natalucci L., Roques J. P., 2008, in Proceedings of the 7th INTEGRAL Workshop, p. 144
- Mizuno T., Kamae T., Godfrey G., Handa T., Thompson D. J., Lauben D., Fukazawa Y., Ozaki M., 2004, ApJ, 614, 1113
- Penacchioni A., Braga J., Castro M., D’Amico F., 2015, J. High Energy Astrophys., 5, 22
- Pottschmidt K., Chernyakova M., Zdziarski A. A., Lubiński P., Smith D. M., Bezayiff N., 2006, A&A, 452, 285
- Remillard R. A., McClintock J. E., 2006, Annu. Rev. Astron. Astrophys., 44, 49
- Sarkar R., Mandal S., Debnath D., Kotoch T. B., Nandi A., Rao A. R., Chakrabarti S. K., 2011, Exp. Astron., 29, 85
- The National Physical Laboratory, 2008, available at: www.kayelab.npl.co.uk

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.