



# A correlation study regarding the AE index and ACE solar wind data for Alfvénic intervals using wavelet decomposition and reconstruction

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**Abstract.** The purpose of this study is to present a wavelet interactive filtering and reconstruction technique and apply this to the solar wind magnetic field components detected at the L1 Lagrange point  $\sim 0.01$  AU upstream of the Earth. These filtered interplanetary magnetic field (IMF) data are fed into a model to calculate a time series which we call AE\*. This model was adjusted assuming that magnetic reconnection associated with southward-directed IMF  $B_z$  is the main mechanism transferring energy into the magnetosphere. The calculated AE\* was compared to the observed AE (auroral electrojet) index using cross-correlation analysis. The results show correlations as high as 0.90. Empirical removal of the high-frequency, short-wavelength Alfvénic component in the IMF by wavelet decomposition is shown to dramatically improve the correlation between AE\* and the observed AE index. It is envisioned that this AE\* can be used as the main input for a model to forecast relativistic electrons in the Earth's outer radiation belts, which are delayed by  $\sim 1$  to 2 days from intense AE events.

## 1 Introduction

Around solar maximum the major causes of geomagnetic storms and space weather disturbances on Earth are interplanetary coronal mass ejections (ICMEs), especially magnetic clouds (MCs), and their sheath or shocked fields (Gonzalez et al., 2007; Echer et al., 2011). However, during the descending and minimum solar cycle phases, high-speed solar

wind streams (HSSs) become the dominant interplanetary–heliospheric structure, causing geomagnetic activity on Earth (Sheeley et al., 1976; Tsurutani and Gonzalez, 1987; Tsurutani et al., 1995, 2006; Guarnieri, 2006; Guarnieri et al., 2006; Kozyra et al., 2006; Turner et al., 2006; Gonzalez et al., 2007; Echer et al., 2011; Hajra et al., 2013). The features within the HSS causing geomagnetic activity are large-amplitude Alfvén waves (Belcher and Davis, 1971; Tsurutani et al., 1982, 1990, 2011a, b; Echer et al., 2011; Hajra et al., 2013), the southward component of which leads to intermittent magnetic reconnection between the wave magnetic fields and the magnetopause fields, allowing the transfer of energy and momentum from the solar wind to the magnetosphere (Dungey, 1961).

One of the most widely used indices to estimate the energy input to the magnetosphere and ionosphere is the geomagnetic auroral electrojet (AE) index. The index is the maximum deviation of the horizontal components of geomagnetic field variations from a set of globally distributed ground-based magnetometers located in and near the auroral zone in the Northern Hemisphere. The AE index represents the overall disturbances in both eastward and westward ionospheric electrojets located at  $\sim 100$  km altitude (Sugiura and Davis, 1966). Thus, substorms and injection events causing magnetotail plasmashet injections into the midnight sector of the magnetosphere with concomitant particle precipitation into the auroral zone ionosphere may intensify both electrojets, leading to AE index increases.

The Alfvén waves causing the geomagnetic (AE index) activity have short wavelengths, ranging from  $\sim 2 \times 10^5$  to  $2 \times 10^6$  km ( $\sim 10$  min to  $\sim 2$  h in duration convected in a  $400 \text{ km s}^{-1}$  solar wind), much smaller than ICME and MC scales. The direct use of the IMF  $B_z$  at the L1 libration point during such structures results in poor correlation against the AE on Earth (Tsurutani et al., 1990, 1995, Guarnieri, 2005). It is thought that part of the problem is that Alfvén waves are propagating in the solar wind, and what is detected at the L1 point is not what hits the Earth's magnetosphere. Another possibility is that the high-frequency wave power does not contribute to the solar wind energy transfer process.

A second incentive for better understanding the relation between interplanetary structures and the AE indices is that intense AE activity events called HILDCAAs (Tsurutani and Gonzalez, 1987) have been shown to be indirectly related to the production of relativistic electrons in the Earth's magnetosphere (Hajra et al., 2014, 2015a, b). In particular Hajra (2015a) showed that HILDCAA onsets precede relativistic  $\sim 0.6$  MeV electron acceleration by  $\sim 1$  day and  $\sim 4.0$  MeV electron acceleration by  $\sim 2$  days. These HILDCAA events are well correlated with the presence of Alfvén waves within HSS (Tsurutani and Gonzalez, 1987; Gonzalez et al., 1994), and the intensity of the geomagnetic event depends on the amplitude of the negative  $B_z$  component of the magnetic field of these waves (Guarnieri, 2005, 2006; Guarnieri et al., 2006).

The overall picture of relativistic electron acceleration in the magnetosphere is the following: reconnection between the southward component of the Alfvén waves and the Earth's dayside magnetopause field (Dungey, 1961; Gonzalez and Mozer, 1974; Tsurutani et al., 1995) leads to substorms and convection events and injections of energetic electrons into the nightside region of the outer magnetosphere (DeForest and McIlwain, 1971; Horne and Thorne, 1998). The energetic electron component creates electromagnetic chorus waves through the loss cone instability (Tsurutani and Smith, 1977; Meredith et al., 2001; Tsurutani et al., 2013). Then the chorus accelerates the high-energy electrons to relativistic energies by resonant interactions (Inan et al., 1978; Horne and Thorne, 1998; Thorne et al., 2005, 2013; Summers et al., 2007; Reeves et al., 2013; Boyd et al., 2014; Hajra et al., 2015a).

The acceleration of relativistic electrons within the Earth's outer radiation belt (Paulikas and Blake, 1979; Baker et al., 1986) is an important physical phenomenon in space weather. These electrons are also known as “killer electrons” for their hazardous effects to orbiting spacecraft (Wrenn, 1995; Horne, 2003). Recent studies by Hajra et al. (2013, 2014, 2015a, b) indicate the probability that magnetospheric relativistic electron acceleration may be predicted more than 1 day in advance using ground-based observations of auroral activities during HSS.

This paper describes a correlation analysis using a technique of wavelet decomposition and selective reconstruction

applied in both IMF solar wind data and AE index. The results of this technique may allow us, in the future, to develop a more complete model to forecast the occurrence of relativistic electrons during periods with Alfvénic fluctuations in the interplanetary solar wind.

## 2 Methodology

Interplanetary magnetic field and solar wind parameters obtained from the ACE (Advanced Composition Explorer) spacecraft (Stone et al., 1998) were used in this work. This data set has  $\sim 1$  min resolution, and we have used the level 2 processed data. IMF vector data used in this work are in the GSM coordinate system. The ACE spacecraft is located at the L1 libration point,  $\sim 1.5$  million km from the Earth, orbiting a region around the Sun–Earth line. The data are available online at <http://www.srl.caltech.edu/ace>.

The geomagnetic activity was observed through the AE indices (Sugiura and Davis, 1966) and Dst index (Rostoker, 1972). These indices are available through the World Data Center for Geomagnetism, Kyoto (<http://wdc.kugi.kyoto-u.ac.jp>). The AE and Dst indices have 1 min and 1 h time resolutions, respectively. The AE index was used to identify the periods of enhanced auroral electrojet activity, while the Dst index was used only to ensure that the analysed intervals were not occurring during the main phases of magnetic storms.

For this study, we used a set of 14 geoeffective interplanetary HSS events, previously identified by Guarnieri (2005) as the longest-lasting elevated AE index events spanning 1998–2001. Table 1 shows a listing of these events with the year, event start date and time, event end date and time, and event duration (in minutes). The Alfvénicity of the solar wind for these intervals was verified using the classical technique proposed by Belcher and Davis (1971) (i.e. these elevated AE intervals are associated with high-speed stream solar wind origin).

A filtering process adapted from the Meyer wavelet decomposition and reconstruction was employed in this technique. This procedure allows a decomposition of the signal into bands with periods in multiples of  $2^n$  of the data cadence (1 min), with  $n = 1, 2, 3, \dots$ . Each decomposed band is called a “detail” and represented by  $D_n$ , where  $n$  represents the decomposition level. Table 2 shows each decomposition level  $D_n$ , the level  $n$ , the associated period ( $2^n$ ), and the period range in minutes. More details about this technique can be found in Meyer (1993) and Kumar and Fofoula-Georgiou (1997).

The last level indicated in Table 2 (A10) contains all the periodicities longer than 1024 min ( $\sim 17.1$  h), and it can be considered as the residual of the decomposition process. Further, this level contains the average value of the data series. We choose this level to terminate the decomposition since de-

**Table 1.** Long-lasting AE events occurred between years 1998 and 2001 (Guarnieri, 2005).

Year	Event	Start	End	Duration (min)
		Date and time (UT)	Date and time (UT)	
1998	1	24 Apr 18:03	27 Apr 06:05	3603
	2	22 Jul 21:09	25 Jul 12:25	3797
1999	1	29 Apr 11:20	3 May 11:16	5757
	2	17 Aug 22:52	20 Aug 12:00	3669
	3	31 Aug 15:32	2 Sep 20:30	3179
	4	10 Oct 20:00	14 Oct 17:38	5619
	5	23 Oct 13:21	25 Oct 20:57	3337
	6	7 Nov 17:00	10 Nov 04:47	3588
	7	3 Dec 10:00	6 Dec 00:15	3736
2000	1	27 Jan 18:10	31 Jan 03:15	4866
	2	5 Feb 16:01	8 Feb 05:33	3693
	3	24 Feb 00:03	27 Feb 22:10	5648
	4	24 May 10:00	26 May 18:07	3368
2001	1	11 May 14:04	14 May 10:51	4128

**Table 2.** Decomposition levels and the corresponding periods range for each level used in the wavelet decomposition technique.

Level	$n$	Range (minutes)
D1	1	< 2
D2	2	2–4
D3	3	4–8
D4	4	8–16
D5	5	16–32
D6	6	32–64
D7	7	64–128
D8	8	128–256
D9	9	256–512
D10	10	512–1024
A10	–	> 1024

tails of higher orders are so smoothed that they are not useful for the AE evaluation.

If two or more details are taken and added the time series, this will result in an “approximation”, which can be considered as a band pass filter. Taking the A10 level and adding it to D10 will result in approximation A9, and so on ( $A_{n-1} = A_n + D_n$ ). In this way, the A0 level will be exactly the same as the original signal, since it contains all the decomposition levels. The reconstruction is an interactive process that can be started and stopped at any decomposition level.

A computer routine was developed to adjust the parameters of the empirical equations for the calculated AE. After the adjustments, the model was fed with the filtered IMF and solar wind data time series. The calculated AE for each

event was compared to the real geomagnetic index observed to check the correlation among them.

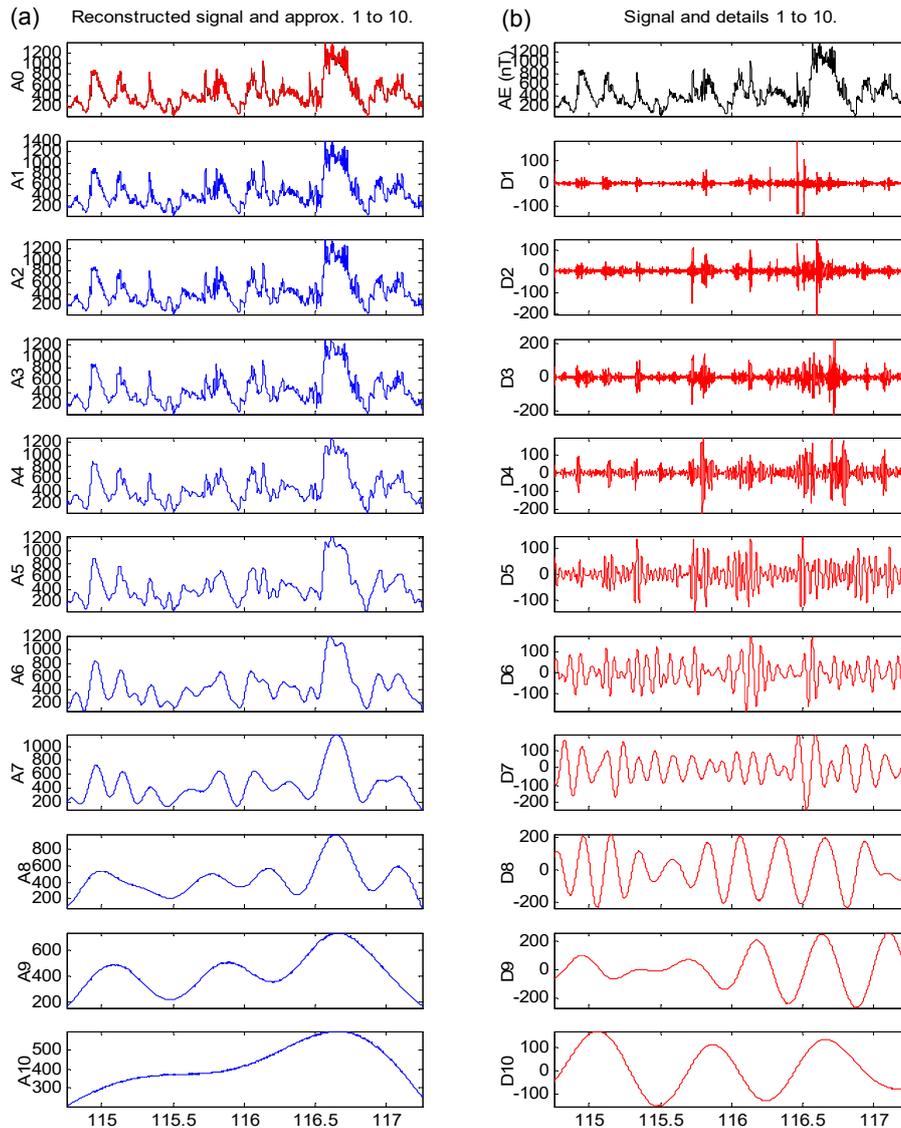
### 3 Results and discussion

Preliminary tests were performed using the cross-correlation analyses between the AE index and several interplanetary parameters, such as  $|B|$ ,  $B_x$ ,  $B_y$ ,  $B_z$ ,  $V_{sw}$ , and  $N_p$ , as was done previously by Guarnieri (2005). The  $B_z$  magnetic field component was found to be the parameter most related to the auroral activity. For this reason, the attempts to adjust a function describing the AE index were mostly focused on this IMF component.

The wavelet decomposition technique was applied in both the AE index and the IMF  $B_z$  component. Figure 1 shows an example of wavelet decomposition and reconstruction for the AE index. The top right panel shows the AE time series. The panels in the right side are the “details”, identified by D1 to D10 (see Table 2 for the corresponding range of each detail). Periods longer than 1024 min plus the average value of the data series are in the A10 level, shown in the bottom panel, left side. The panels in the left side are the “approximations”, which can be viewed as a cumulative sum of details.

Since the high-speed solar wind events are characterized by enhanced AE activity, the higher approximation levels (such as A8, A9, and A10) present increasingly averaged values.

Figure 2 shows the wavelet decomposition and reconstruction for the solar wind magnetic field  $B_z$  component. The sequence of panels is the same as Fig. 1. Again, the A10 level represents a typical characteristic during high-speed solar wind streams: a small but continuous negative  $B_z$  average value. This characteristic was already observed by



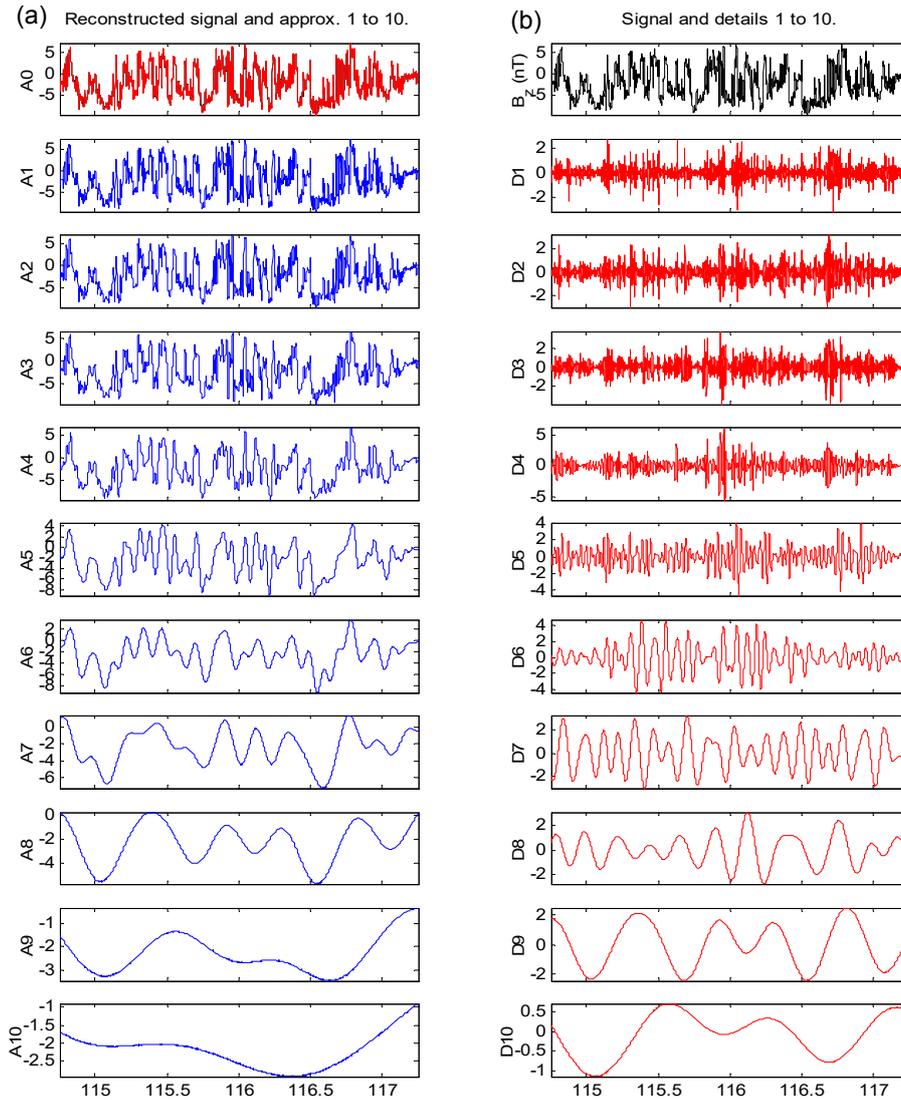
**Figure 1.** Wavelet decomposition of the AE index for event 1\_1998. The right side shows the “details” (bands) and the left side shows the “approximations” (reconstructions).

Guarnieri (2005), through the calculation of average values of  $B_z$  during HSS intervals.

Comparing Figs. 1 and 2, an anti-correlation in level A10 is clearly observable. This same behaviour is also present in other approximation levels. This anti-correlation shows that AE activity is driven by the  $-B_z$  on these long timescales during the Alfvénic intervals.

Previous work (Guarnieri, 2005) had observed that high frequencies in the signal could hide or decrease the correlation between solar wind parameters and geomagnetic indices, because of the presence of noisy, turbulent activity. So, an approximation level has to be chosen in order to avoid these high frequencies and, at the same time, be able to represent the particularities of the signal. With a computer routine,

each reconstruction level was tested and it was found that the correlation is high up to the level A3 (starting from A10). Levels A0, A1, and A2 include most of the high frequencies that reduce the correlation and do not significantly improve the signal characterization. In this work, we used reconstructions from A10 to A3, meaning that only periods longer than 8 min were used in the model. This decision, as well the other assumptions that were used to develop the following empirical equations, were based on several analyses reported in Guarnieri (2005). In that work, Guarnieri used the classical cross-correlation technique, power spectrum, and multi-taper analysis to correlate the  $B_z$  and AE, and the results were only significant when periods shorter than 8 or 16 min (depending on the technique) were removed. The studies were



**Figure 2.** Wavelet decomposition on the interplanetary magnetic field  $B_z$  component for event 1\_1998. The right panels show the details (bands) and the left panels show the approximations (reconstructions).

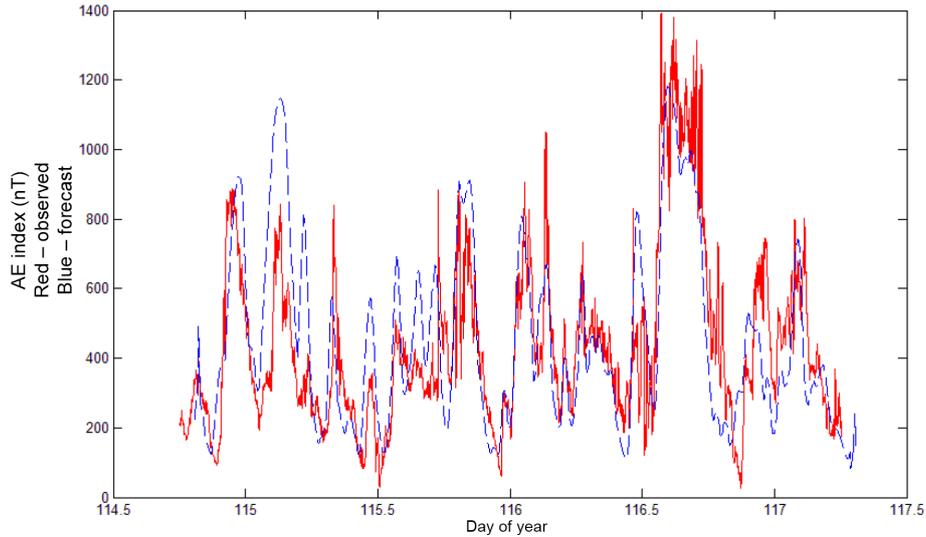
performed in both ACE and IMP-8 data. There was no clear correlation employing unfiltered data and even the lag between the two time series showed inconclusive results. Progressively removing the high frequencies lead to correlation increases and the lag between time series becomes more consistent. The values obtained for correlation were in the range from  $\sim 0.5$  to  $\sim 0.8$ . The lowest values were related to events with the presence of “patches” of different periodicities in  $B_z$ . This behaviour exposes a limitation of the classical correlation techniques in dealing with time-located periodicities.

Once we verified the good correlation between  $B_z$  and AE, an empirical model was developed to estimate a time series (here called AE\*) based on interplanetary  $B_z$  and compare it against the observed AE index.

The first assumption for this empirical model is that reconnection is the main physical process transporting energy from the solar wind into the magnetosphere (and later to the auroral region). In this way, when  $B_z$  is negative we would have energization of the auroral electrojets. Positive  $B_z$  intervals implies that there is no energy input from the solar wind to the magnetosphere, and thus for these intervals a decay function is used to estimate the auroral current decay.

The calculation and modelling process starts with the interplanetary  $B_z$  measured at L1 ( $B_z\_interp$ ), shifted in time to take into account the interplanetary structure travel time from the L1 libration point to the Earth, using the solar wind velocity as a proxy.  $B_z$  is the shifted time series and  $\delta$  is the delay applied:

$$B_z(t) = B_z\_interp_{(t+\delta)}. \quad (1)$$



**Figure 3.** Comparison between the calculated AE\* (dashed blue line) and the real AE index (solid red line) for event 1\_1998.

$B_z$  is then decomposed and reconstructed up to the approximation level desired to eliminate high frequencies, creating the  $B_z^*$  time series (approximation for the shifted  $B_z$ ).

The field change is calculated as follows:

$$\Delta B_{z(t)}^* = B_{z(t)}^* - B_{z(t-1)}^*. \quad (2)$$

The first item of  $ae^*$  is assumed to be  $-B_z^*$  (where  $ae^*$  represents the calculated index before scale and baseline adjustments).

If  $B_{z(t)}^* \leq B_{z(t-1)}^*$ , then the  $B_z$  is getting smaller or more negative, leading to energization:

$$ae_{(t)}^* = ae_{(t-1)}^* + \varepsilon \cdot \Delta B_{z(t)}^*. \quad (3)$$

If  $B_{z(t)}^* > B_{z(t-1)}^*$ , then the following is true:

$$ae_{(t)}^* = ae_{(t-1)}^* \exp(-\gamma B_{z(t)}^*). \quad (4)$$

Finally, the scale and baseline are adjusted in the calculated series:

$$AE^* = \alpha + (\beta \cdot ae^*), \quad (5)$$

where  $AE^*$  is the approximation for the AE index time series. A computational routine was developed to adjust the parameters  $\alpha$ ,  $\beta$ ,  $\varepsilon$ , and  $\gamma$ . The best results were achieved with  $\alpha = 70$  nT,  $\beta = 150$ ,  $\varepsilon = -0.3333$ , and  $\gamma = 1$ . The delay  $\delta$  (in Eq. 1) depends on the solar wind velocity and the shifting method employed. It is in the range of 30–70 min.

Figure 3 shows a comparison between this calculated  $AE^*$  (blue line) and the observed AE index (red line). The reconstructions were created up to the level A4, meaning that only periods longer than 8 min are present.

There is a good correlation between the two series (calculated and observed), although there are still some scale problems. However, some particularities of the real signal were

represented very well by the calculated signal. This event was chosen due to the presence of unambiguous features, such as those occurring at day  $\sim 116$  and the big peak just after day 116.5. These features were used to test the accuracy of the model under unusual conditions.

A comparison among all the events and calculated time series using different approximation levels is shown in Table 3. The data shown in this table are correlation coefficients between the calculated  $AE^*$  and the observed AE index.

Considering the A3 approximation level, all the events have correlation coefficients higher than 0.7. With the exclusion of events ev1\_2000 and ev3\_2000, all of the remaining 12 events have correlation coefficients higher than 0.85. Event 3\_1999 has correlation coefficients larger than 0.958 for all the approximation levels.

Regarding events ev1\_2000 and ev3\_2000, the low correlation coefficients observed led us to reanalyse the data plots to understand what would be the main difference between these events and the remaining ones. The  $B_z$  data for these two events are basically high-frequency oscillations around a  $\sim 0$  nT value, without the longer-period excursions to negative  $B_z$  values typically associated with AE energization. These high frequencies are exactly those mostly removed by the wavelets filtering technique we employed here, leading to a weak correlation against the AE index. Similar results were found by Guarnieri (2005) when analysing these exact two events, but using different techniques, and reaching the same conclusion.

If one tries to apply Eqs. (1) to (5) to unfiltered  $B_z$  data, this may result in a very poor correlation coefficient between the estimated  $AE^*$  and the observed AE index. The interactive filtering process using wavelet decomposition allows us to effectively remove the high-frequency components that

**Table 3.** Correlation coefficients between the calculated (AE\*) and the observed (AE index) series for each event and in each wavelet decomposition approximation level.

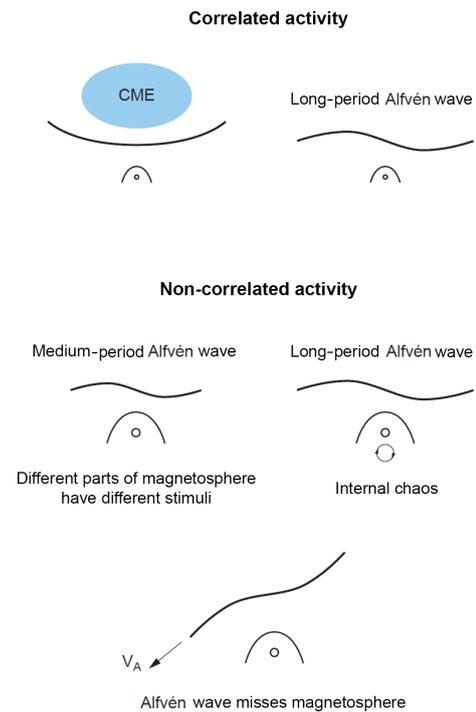
Event	A3	A4	A5	A6	A7	A8	A9
ev1_1998	0.943	0.947	0.948	0.944	0.926	0.928	0.977
ev2_1998	0.913	0.908	0.903	0.907	0.936	0.892	0.894
ev1_1999	0.920	0.920	0.926	0.924	0.931	0.920	0.946
ev2_1999	0.876	0.875	0.865	0.884	0.860	0.853	0.879
ev3_1999	0.958	0.962	0.964	0.965	0.982	0.990	0.995
ev4_1999	0.893	0.889	0.875	0.845	0.844	0.855	0.941
ev5_1999	0.918	0.922	0.915	0.904	0.912	0.901	0.943
ev6_1999	0.921	0.919	0.918	0.902	0.910	0.847	0.863
ev7_1999	0.892	0.887	0.880	0.863	0.857	0.873	0.867
ev1_2000	0.708	0.735	0.866	0.917	0.528	0.542	0.555
ev2_2000	0.856	0.859	0.872	0.891	0.900	0.878	0.890
ev3_2000	0.790	0.597	0.589	0.541	0.443	0.407	0.436
ev4_2000	0.936	0.937	0.936	0.867	0.878	0.906	0.956
ev1_2001	0.908	0.909	0.905	0.906	0.893	0.899	0.925

have poor predictive value, thus obtaining higher correlation values.

These observations lead to several possible scenarios, as shown in Fig. 4. Correlated activity between interplanetary structures and geomagnetic indices are usually related to large interplanetary structures, such as interplanetary coronal mass ejections or long-period Alfvén waves. These structures appear in interplanetary data as low-frequency waves. Due to the sizes of these structures, the Earth’s magnetosphere may react as a whole, and so the geomagnetic indices give us a good idea of the global magnetosphere energization.

However, there are events with uncorrelated activity, usually those with high frequencies present in the interplanetary data, which are related to medium- and short-period Alfvén waves, that may miss impingement on the magnetosphere. There are also long-period waves that can be uncorrelated due to internal chaotic processes inside the magnetosphere. When these high-frequency events are removed by filtering IMF  $B_z$  and AE index data, we are able to reach high correlation values such as those shown on Table 3.

Future work can implement a forecasting model for the AE index for such periods with high-amplitude Alfvénic fluctuations. One has to use the real-time data from a spacecraft around L1 and verify the Alfvénicity through the technique employed by Belcher and Davis (1971). Once the Alfvénicity is verified, the data series would be fed in through the wavelet filtering and the AE\* evaluation equations. This would give us a forecasted AE with a delay of only a few minutes, or almost a “nowcast”. However, the main result would be the relativistic electron forecasting. Since Hajra et al. (2013, 2014, 2015a, b) have shown the probability that magnetospheric relativistic electron acceleration may be predicted by more than 1 day in advance using ground-based observations of auroral activities during high-speed solar wind streams, the ability to obtain an equivalent auroral index may lead to a



**Figure 4.** Schematic showing the possible causes for correlated and uncorrelated events between interplanetary parameters and geomagnetic indices.

more complete model which also includes the forecasting of relativistic electrons. It is important to note that data from other spacecraft or different processing levels can be employed in this forecast. One just has to take in to account the propagation delay according to the spacecraft position.

## 4 Conclusions

A method to calculate an AE\* time series based on IMF data measured at L1 in the presence of interplanetary Alfvén waves was developed. This method employs an Alfvénicity check and a wavelet decomposition technique, which is applied to the interplanetary magnetic field  $B_z$  component. This calculated AE\* was shown to be highly correlated with the observed AE index. The correlation coefficients between the calculated and observed series can reach values over 0.90, depending on the resolution and the level of details assumed.

Future works can use data from the any other spacecraft located around L1 to feed the model and obtain the AE\*, and this AE\* feeds a routine to predict the occurrence of relativistic electrons, giving advanced notice by more than a day. This work will hopefully be completed within the next few years.

## 5 Final comments

Although correlations between AE\* and AE as high as 0.90 have been indicated in this paper, the causes for the lack of a high correlation in some intervals are still in question. One possibility is that there is high-frequency turbulence in the solar wind which is not geoeffective, a point mentioned previously. However there is another, recently discussed possibility: local generation of Alfvén waves in the interplanetary medium (Tsurutani et al., 2018). If Alfvén waves are generated between ACE and the Earth, this would naturally reduce the correlation between AE\* and AE. This may account for intervals where AE\* and AE are not well correlated. Both of these explanations are possible. In the future we will probe which of these two (or other possibilities) are factors, and which are the dominant ones.

*Data availability.* ACE spacecraft data is available at <http://www.srl.caltech.edu/ace> (ACE, 2007). Geomagnetic data (AE and Dst indices) are available at <http://wdc.kugi.kyoto-u.ac.jp/> (World Data Center for Geomagnetism, 2007).

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